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ARE ELIAS 2-27'S SPIRAL ARMS DRIVEN BY SELF-GRAVITY, OR BY A COMPANION?
A COMPARATIVE SPIRAL MORPHOLOGY STUDY

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(Received January 1, 2018; Revised January 1, 2018; Accepted January 1, 2018)

Submitted to ApJ

ABSTRACT

The spiral waves detected in the protostellar disc surrounding Elias 2-27 have been suggested as evidence of the disc being gravitationally unstable. However, previous work has shown that a massive, stable disc undergoing an encounter with a massive companion are also consistent with the observations. We compare the spiral morphology of smoothed particle hydrodynamic simulations modelling both cases. The gravitationally unstable disc produces symmetric, tightly wound spiral arms with constant pitch angle, as predicted by the literature. The companion disc's arms are asymmetric, with pitch angles that increase with radius. However, these arms are not well-fitted by standard analytic expressions, due to the high disc mass and relatively low companion mass. We note that differences (or indeed similarities) in morphology between pairs of spirals is a crucial discriminant between scenarios for Elias 2-27, and hence future studies *must* fit spiral arms individually. If Elias 2-27 continues to show symmetric tightly wound spiral arms in future observations, then we posit that it is the first observed example of a gravitationally unstable protostellar disc.

Keywords: stars: individual (Elias 2-27) — stars: pre-main sequence — hydrodynamics — protoplanetary disks — planet-disk interactions

1. INTRODUCTION

Spiral structures generated by gravitationally unstable protostellar discs play a crucial role in the evolution of protostars and the planetary systems they eventually host. At the instant of a system's formation, the star mass and disc mass are comparable. This guarantees that the Toomre Parameter (Toomre 1964):

$$Q = \frac{c_s \kappa_{\text{epi}}}{\pi G \Sigma} \sim 1, \quad (1)$$

where c_s is the sound speed of the disc gas, κ_{epi} is the epicyclic frequency (which is equal to the angular frequency Ω if the disc is Keplerian) and Σ is the surface density. As such, discs which satisfy $Q \sim 1$ will be unsta-

ble to non-axisymmetric perturbations, which undergo swing amplification into spiral structures.

At early times, these structures are typically strong global modes, which can achieve rapid accretion of the disc onto the star via non-local angular momentum transport (Laughlin & Bodenheimer 1994; Lodato & Rice 2005; Forgan et al. 2011). Under the appropriate conditions, spiral arms can fragment into gravitationally bound objects, representing a formation channel for low mass stars, brown dwarfs, gas giant planets and in some very rare cases, terrestrial planets (Gammie 2001; Rice et al. 2005; Stamatellos et al. 2007; Galvagni & Mayer 2014; Forgan & Rice 2013a; Forgan et al. 2018a).

Constraining both disc fragmentation as a formation mechanism, and protostellar accretion in general, requires us to observe *bona fide* gravitationally unstable protostellar discs in the wild. For observational campaigns, there are (broadly) two approaches to determin-

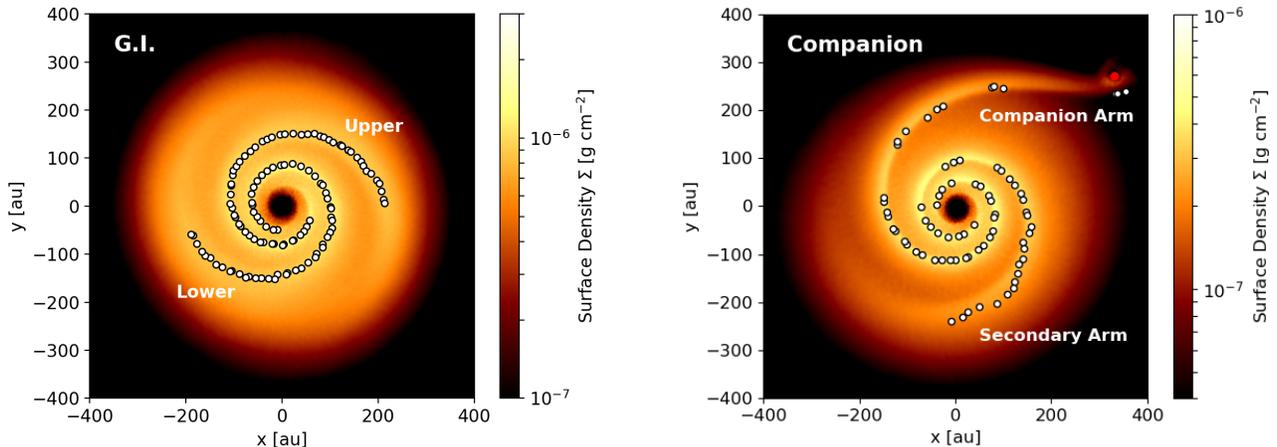


Figure 1. Surface density of the gravitationally unstable disc (left) and the companion encounter disc (right) overlaid with the results of the TACHE spiral identification algorithm (black crosses). Individual arms are identified as either ‘upper’ or ‘lower’ in each case.

ing whether a protostellar disc is gravitationally unstable - measuring its physical properties and determining Q , or studying its morphology for signs of spiral structure.

Determining the physical properties of self-gravitating protostellar discs is a challenging endeavour, as the self-gravitating phase is brief due to rapid disc accretion. As a result, self-gravitating discs remain heavily embedded inside their birth molecular cloud. Further to this, their centrally condensed surface density profiles can frustrate efforts to measure disc masses through measuring continuum dust emission, due to large optical depths even at sub-millimetre wavelengths (Greaves & Rice 2010; Forgan & Rice 2013b; Forgan et al. 2016b; Evans et al. 2017).

The detection and characterisation of spiral structure, then, may yield a bias-free determination of whether a protostellar disc is gravitationally unstable. Spirals in protostellar discs have recently come within the reach of observations. Scattered light measurements have yielded several protostellar discs exhibiting spirals (e.g. Fukagawa et al. 2006; Muto et al. 2012a; Benisty et al. 2015; Dong et al. 2016), but these observations only trace structure generated in the disc’s upper layers. Gravitationally unstable discs will drive structure in the bulk of the disc material, at all altitudes, down to the midplane.

The arrival of the Atacama Large Millimetre Array (ALMA) has allowed exquisite spatially resolved observations of protostellar discs at sufficiently long wavelengths to probe spiral structure at the disc midplane. A most striking recent example is Elias 2-27 – a relatively young low mass star ($M_* \sim 0.5 - 0.6M_\odot$, age \sim

1 Myr; Luhman & Rieke 1999; Natta et al. 2006), hosting a Class II circumstellar disc potentially exhibiting a large mass ($M_{\text{disc}} \sim 0.04 - 0.14M_\odot$, Andrews et al. 2009; Isella et al. 2009; Ricci et al. 2010).

Pérez et al. (2016) presented ALMA observations that showed two large-scale, symmetric spiral arms. In their study, both arms were simultaneously fitted to pure logarithmic spirals, with identical pitch angle $7.9^\circ \pm 0.4^\circ$. Given both its relatively large disc-to-star mass ratio ($q \sim 0.06 - 0.3$) and the presence of spiral structure driven at the midplane, it has been suggested that Elias 2-27 is a gravitationally unstable disc system (Tomida et al. 2017; Meru et al. 2017).

Of course, gravitational instability (GI) is not the only mechanism that generates spiral structures. Interactions between a disc and a companion generate tidally-driven arms, that (to the eye) can be very similar to arms driven in an isolated GI disc. Meru et al. (2017) address this issue by running a suite of smoothed particle hydrodynamics (SPH) simulations of both isolated GI discs, and GI discs that are perturbed by a companion. They show that both isolated and perturbed discs produce spiral structures that, when observed synthetically with ALMA, using the same unsharp masking technique as Pérez et al. (2016), produce images consistent with Elias 2-27’s features.

Hall et al. (2018) also consider a range of synthetically observed isolated GI disc simulations, and show that if Elias 2-27 is an isolated GI disc, then its properties are tightly constrained, where slight changes to its physical properties result either in dissipation of the spirals or fragmentation. This suggests that (*a priori*) Elias 2-27 is less likely to be an isolated GI disc.

We therefore argue that comparative morphology studies are a crucial orthogonal tool to determine whether GI or a companion is driving spiral structure in a given protostellar disc. In this work, we perform such a study on the simulations of Meru et al. (2017) to identify crucial differences in spiral morphologies between isolated and perturbed GI discs. Our results offer several discriminants for determining the nature and origin of spiral structure in massive protostellar discs.

2. METHODOLOGY

2.1. Hydrodynamics

Our hydrodynamic simulations are fully described in Meru et al. (2017), but for completeness we briefly reiterate some salient aspects. The simulations are performed using the three-dimensional Smoothed Particle Hydrodynamics code (SPHNG) including heating due to work done and the radiative transfer of energy in the flux-limited diffusion limit (Whitehouse et al. 2005; Whitehouse & Bate 2006). A detailed description of the code can be found in Meru (2015) and Meru et al. (2017).

Our reference model consists of a $0.5M_{\odot}$ star, modelled as a sink particle, surrounded by a disc whose initial temperature (T) and surface density (Σ) vary with radius (R) as

$$T(R) = 13.4 \text{ K} \left(\frac{R}{200 \text{ au}} \right)^{-0.75}, \text{ and} \quad (2)$$

$$\Sigma(R) = \Sigma_0 \left(\frac{R}{200 \text{ au}} \right)^{-0.75} \quad (3)$$

respectively, between $R_{\text{in}} = 10 \text{ au}$ and $R_{\text{out}} = 350 \text{ au}$. The first simulation is a gravitationally unstable disc (GI), with $\Sigma_0 = 6 \times 10^{-7} M_{\odot} \text{ au}^{-2}$, giving a total disc mass of $0.24 M_{\odot}$, and an initial Toomre parameter $Q < 1$ beyond $\sim 250 \text{ au}$.

In the second simulation (Companion), $\Sigma_0 = 1.96 \times 10^{-7} M_{\odot} \text{ au}^{-2}$, giving a total disc mass of $0.078 M_{\odot}$, and $Q > 2$ at all radii.

We note that for the GI case the disc does evolve, such that $\Sigma \propto R^{-0.5}$. For the Companion simulation we model an $8M_{\text{Jup}}$ companion initially located at 500 au that is allowed to freely interact with the disc, migrate and grow. At the time when the simulation is analysed the companion is $\approx 10M_{\text{Jup}}$ and located at $\approx 425 \text{ au}$ from the central star. The companion does not drive a gap, which is consistent with the gap opening criteria defined by both Lin & Papaloizou (1986) and Crida et al. (2006). We note that the presence or absence of a gap makes little difference to our results. Each disc is modelled using 250,000 SPH gas particles, and we assume that the gas and dust are well mixed (see Meru et al. 2017, for further details).

2.2. Spiral Detection & Morphology

We use the TACHE code, which utilises tensor-classification of the simulations to determine which SPH particles reside in spiral structures (Forgan et al. 2016a, 2018b). Briefly, we compute the velocity shear tensor of each particle:

$$\sigma_{ij} = -\frac{1}{2} \left(\frac{\partial v_i}{\partial x_j} + \frac{\partial v_j}{\partial x_i} \right). \quad (4)$$

And then compute the tensor's eigenvalues. The number of positive eigenvalues E encodes information about the dimensionality of the flow. For example, particles with $E = 1$ indicate their motion is planar, consistent with motion in the undisturbed disc. Particles with $E = 2$ indicate 2D filamentary motions (in this case, spiral structure). If the spiral structure is strong enough, particles near the centre of the arm will possess $E = 3$ (3D collapse).

We can therefore identify particles belonging to the spiral using their E value, and discard the other particles. This allows us to trace the spine of the spiral structure (i.e. the location of maximum density), using a friends-of-friends-like algorithm, which yields a set of (x, y) points for each individual spiral. Each arm is then fitted separately via χ^2 minimisation to a variety of spiral models (assuming a constant uncertainty of $\sigma = 0.1 \text{ AU}$ for all points). We use Nelder-Mead (amoeba) optimisation to obtain said minimum χ^2 , implemented via `scipy.optimize.minimize`.

Logarithmic spirals are commonly found in isolated GI discs and in discs driven by encounters with a companion:

$$r = ae^{b\theta}, \quad (5)$$

where r is cylindrical radius, and θ is the azimuthal angle. The a parameter determines the initial distance of the spiral from the origin, and b determines the winding properties of the arm. The pitch angle of a logarithmic spiral

$$\phi = \left| r \frac{d\theta}{dr} \right|^{-1} = \arctan b. \quad (6)$$

Pure logarithmic spirals (where $b = b_0$ is a constant) are typically found in simulations of isolated GI discs, with a constant $\phi \sim 10 - 15^\circ$ (for $q < 0.5$ Cossins et al. 2009; Hall et al. 2016; Forgan et al. 2018b). We will label model fits of this type as PURELOG.

We also consider models where the pitch angle varies with radius, as is expected if the spiral is being driven by a companion (Goodman & Rafikov 2001; Rafikov 2002; Muto et al. 2012b). For low mass companions ($M \lesssim 1M_{\text{Jup}}$ in low mass discs ($M_d/M_* \lesssim 0.1$), logarithmic

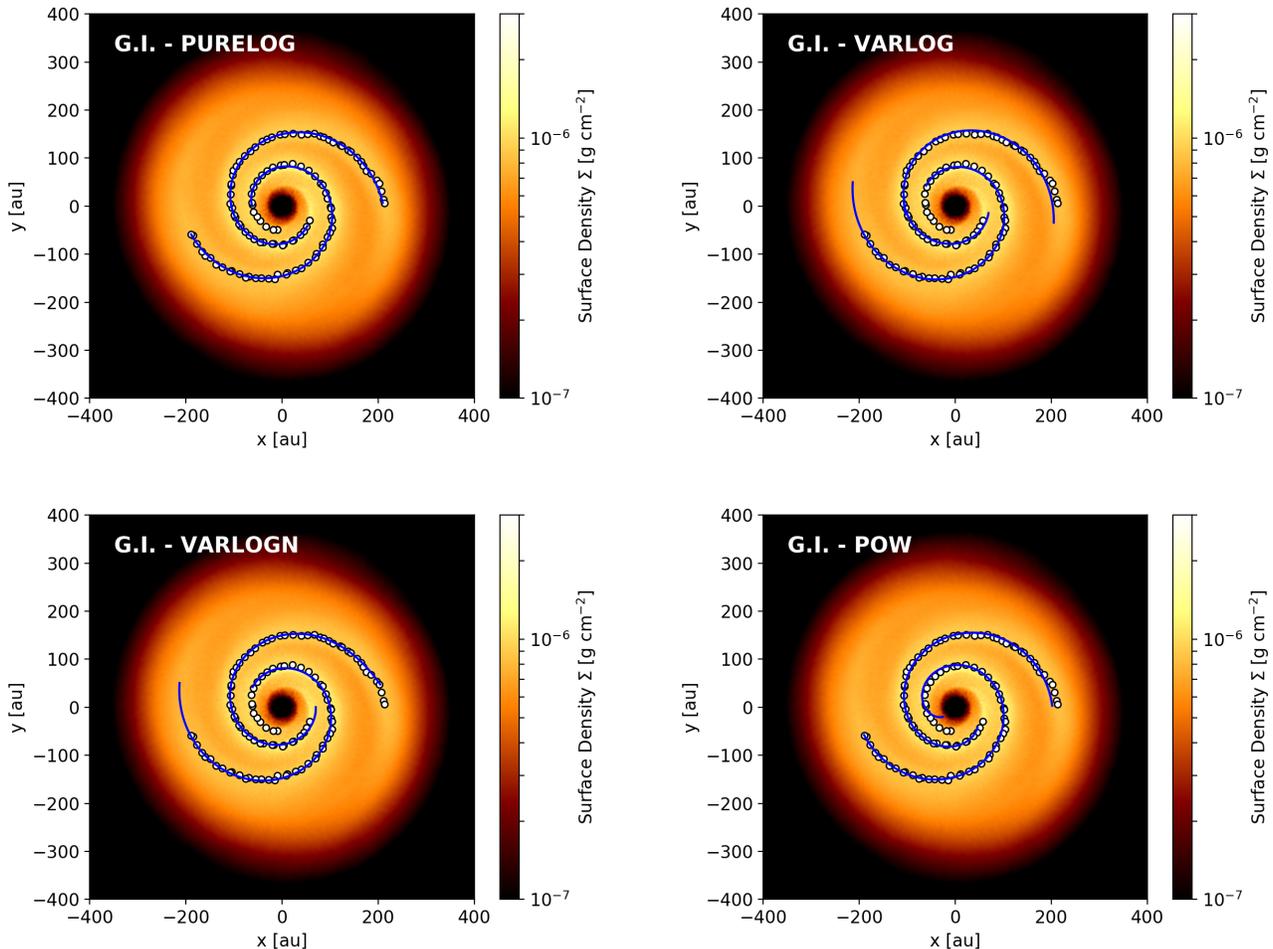


Figure 2. The isolated GI disc simulation, overlaid with the TACHE spine points and four different fitted spiral functions (see text).

spirals are typically found with the following function for b (Zhu et al. 2015),

$$b(r) = h_p \left(\frac{r_p}{r} \right)^{1+\eta} \frac{r^\alpha}{r^\alpha - r_p^\alpha}, \quad (7)$$

where r_p is the orbital radius of the companion $\Omega \propto r^{-\alpha}$, $c_s \propto r^{-\eta}$, and h_p is the aspect ratio of the disc at the location of the companion.

We find that this function gives a very poor fit for both cases, as the spiral arms are deeply non-linear, due to the massive, self-gravitating nature of both discs (and the relatively high mass companion). Despite the lack of theoretical guidance in this regime, we still consider the possibility that the pitch angle does vary with radius. Instead, we fit a simpler function that describes a pitch angle that increases relatively slowly with radius (VARLOG),

$$b(r) = b_0 + c \left(\frac{r-a}{a} \right)^n, \quad (8)$$

where $b = b_0$ at $r = a$, and the PURELOG solution is recovered for $c = 0$. We also check specifically whether a linear dependence of b with r is sufficient by running fits with $n = 1$ (VARLOGN). Finally, to double check that the spirals do indeed possess a logarithmic form, we also fit a power spiral function (POW), where

$$r = a\theta^n \quad (9)$$

which also includes a radius-dependent pitch angle of the form

$$\phi = n \left(\frac{r}{a} \right)^{-1/n}, \quad (10)$$

where for the special cases of $n = \pm 1$, the power spiral reduces to the Archimedean and hyperbolic spirals respectively (modulo constants describing the spiral origin).

3. RESULTS & DISCUSSION

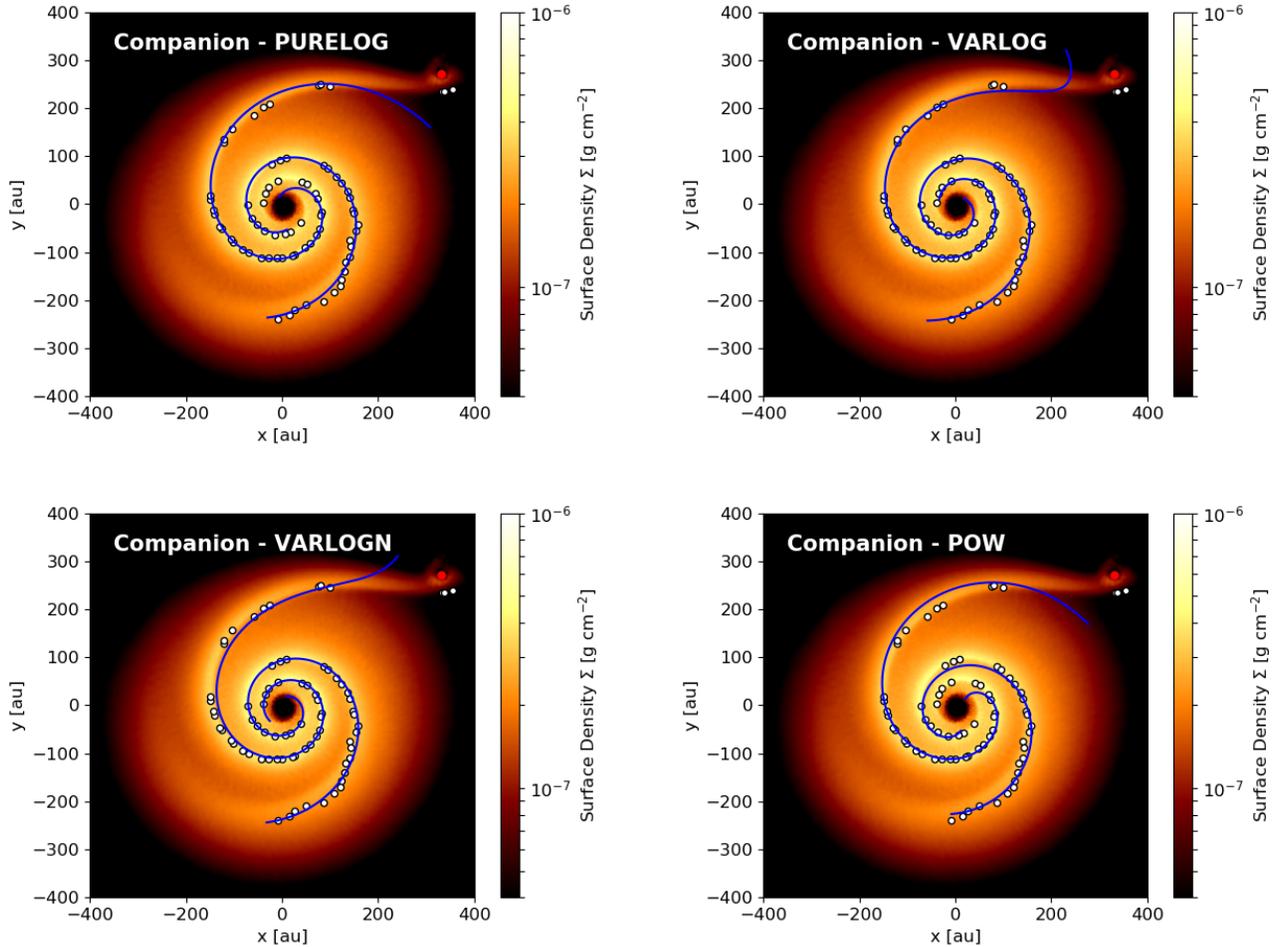


Figure 3. As Figure 2, but for the Companion simulation. The location of the perturbing companion is shown in each panel by a red circle.

Figure 1 shows the results of TACHE’s spiral spine identification for the GI disc (left) and the Companion disc (right). Note that for the Companion disc, the spine identification slightly misses the location of the companion. This is due to the fact that the particles immediately surrounding the companion have a different tensor class to that of the spirals, and are thus removed from the spiral spine fitting.

In the following sections we attempt to fit these spiral spines using the spiral functions described previously. A summary of our best fitting parameters for each case, along with the respective goodness-of-fit statistics, are given in Table 1.

3.1. Isolated GI disc

We find that for the isolated GI disc, pure logarithmic spirals of constant pitch angle (PURELOG) deliver a very good fit to the data. Other spiral functions yield reasonable fits to the arms (Figure 2), but all yield

poorer χ^2 values than PURELOG for the upper arm, and deliver little change on the lower arm. Power spiral fits (POW) produce quite poor fits at the inner and outer regions, strongly indicating that the arm’s pitch angles are indeed constant with radius, as is the case for a PURELOG spiral. VARLOG/VARLOGN fits require extra parameters and fail to yield better fits, again indicating that a constant pitch angle is the simplest and most effective model fit to these arms.

Such a result is in accordance with our expectation that spiral arms in GI discs are density waves, and should hence be pure logarithmic spirals (Cossins et al. 2009; Hall et al. 2016). It is also in accord with TACHE’s previous applications to GI discs with disc-to-star mass ratios less than 0.5 (Forgan et al. 2018b). Note that the derived pitch angles for each arm are within 0.2° (1.4%) of each other. A hallmark of isolated GI discs are arms with extremely similar (if not identical) pitch angles.

Table 1. Best fitting parameters for both the G.I. and Companion simulation for the spiral functions described in Section 2.2. In each case, ‘up’ and ‘low’ refer to the upper and lower arms identified in Figure 1, respectively.

	Gravitational Instability								Companion							
	PURELOG		VARLOG		VARLOGN		POW		PURELOG		VARLOG		VARLOGN		POW	
	Up	Low	Up	Low	Up	Low	Up	Low	Up	Low	Up	Low	Up	Low	Up	Low
a	29.2	56.9	111.8	56.2	110.3	56.8	55.3	11.5	64.2	33.0	39.4	73.3	46.5	38.5	8.4	1.9
b_0	0.208	0.205	0.191	0.244	0.123	0.216	–	–	0.288	0.243	0.208	0.225	0.148	0.201	–	–
ϕ [°]	11.75	11.59	–	–	–	–	–	–	17.0	14.2	–	–	–	–	–	–
c	–	–	-0.009	-0.029	0.1	-0.006	–	–	–	–	0.0003	0.0002	0.014	0.005	–	–
n	–	–	-0.49	0.35	1.0	1.0	0.67	1.26	–	–	2.40	1.92	1.0	1.0	1.64	1.14
χ^2	1.98	3.98	3.63	3.22	3.02	3.26	3.96	4.13	7.80	4.52	3.54	3.85	4.29	2.39	9.91	5.03

A naive conclusion might therefore be that Elias 2-27 is an isolated GI disc, as it too possesses symmetric arms with low pitch angle. However, we must note that companion-driven simulations yield unsharp mask images that also appear to possess symmetric, tightly wound arms (Meru et al. 2017), and that Elias 2-27’s arm symmetry may be an artifact of the fitting procedure (see section 3.3). As we will see in the following section, the “ground truth” morphology of companion-driven spirals is markedly different.

3.2. Companion disc

We can see from Table 1 that PURELOG fits for the companion-driven arms are significantly poorer compared to the VARLOG/VARLOGN fits. The minimum χ^2 solution for PURELOG achieves a reasonable fit at intermediate radii, but at the cost of poorly fitting points in the inner and outer disc. This is heavily indicative that the pitch angle does indeed vary with radius, as is expected for companion-driven arms (Goodman & Rafikov 2001; Rafikov 2002; Muto et al. 2012b; Zhu et al. 2015).

Notably, while the companion-hosting arm is poorly fitted, the secondary arm is much better fitted, with both arms possessing larger pitch angles than the GI run. The companion arm has a pitch angle of almost 17° , which is (just) beyond the $10\text{--}15^\circ$ range expected for isolated GI discs (Cossins et al. 2009). The secondary arm, at $\phi \approx 15^\circ$, therefore appears similar to a typical GI density wave, making it less easy to distinguish.

It is worth noting that analysing subsequent timesteps of the companion simulation yields a range of derived pitch angles for PURELOG, in some cases as large as $\sim 30^\circ$ for the companion arm. We find throughout that the difference in pitch angles between companion and secondary arm are much larger than determined for pairs of arms in the GI disc (section 3.1).

We verify that the spiral form is indeed logarithmic by considering the POW fits. We find that we cannot improve our fits (indeed, they are slightly worse). Again,

we find that the best fit power spiral fails to correctly track the location of the companion.

The best fit to the companion spirals are from the VARLOG/VARLOGN fits. For both fixed n and varying n , we find fits to the companion arm that have χ^2 two to three times smaller than the PURELOG and POW fits. Allowing b to vary constrains the inner spiral much more effectively than the PURELOG fit, and from this analysis it seems clear that the companion arms are indeed logarithmic spirals, not power spirals. However, in contrast to the GI case, the logarithmic spirals are better fitted with a varying pitch angle as opposed to a constant pitch angle.

It is notable that for the VARLOG/VARLOGN fits our functional form for b results in a minimum in $\theta(r)$:

$$\theta(r) = \frac{1}{b(r)} \ln\left(\frac{r}{a}\right), \quad (11)$$

which occurs at

$$\frac{d\theta}{dr} = \frac{1}{br} - \frac{\theta}{b} \frac{db}{dr} = 0. \quad (12)$$

As

$$\frac{db}{dr} = nc \left(\frac{r-a}{a}\right)^n \frac{1}{r-a} = n \frac{b-b_0}{r-a} \quad (13)$$

We therefore find a minimum in $\theta(r)$ at

$$\theta = \left(r \frac{db}{dr}\right)^{-1} = \frac{r-a}{r} \frac{1}{n(b-b_0)}. \quad (14)$$

Hence we see that beyond this minimum, the fitted spiral function turns away from the companion. Note that attempts to fit this function on subsequent timesteps delivers a good fit to the companion arm, and not the secondary arm. Essentially, the length of the arm determines whether VARLOG/VARLOGN fits will capture the entire arm, or turn away at $\frac{d\theta}{dr} = 0$.

This is an important issue for attempts to fit spiral structure driven by companions. Despite TACHE being able to roughly locate the companion inside the spiral

structure, attempts to fit the structure with typical spiral functions uniformly fail to locate the companion correctly. We therefore urge caution when attempting to determine the location of an unseen companion using spiral structure alone. If the companion mass is much less than ~ 0.02 times the stellar mass, the two arms can be significantly asymmetric, and the companion location can be determined using the two spiral arms alone (Fung & Dong 2015).

3.3. Implications for Elias 2-27

Both the GI and Companion simulations, when observed synthetically and subjected to the same unsharp mask imaging techniques as carried out by Pérez et al. (2016), broadly reproduce the observed spiral morphology of Elias 2-27 (Meru et al. 2017). However, our results identify key discriminators between the “ground-truth” spiral morphologies of the two cases.

If Elias 2-27 is an isolated, gravitationally unstable protostellar disc, our simulations predict symmetric, logarithmic spiral arms of constant pitch $\phi \sim 12^\circ$, consistent with expectations from density wave theory (Lin & Shu 1964; Bertin & Lin 1996), which are appropriate as for this case the disc-to-star mass ratio $q \lesssim 0.5$ (Cossins et al. 2009; Forgan et al. 2011; Hall et al. 2016). This is slightly larger than the measured pitch of $\sim 8^\circ$ from Pérez et al. (2016), but we should be encouraged by the fact that our simulations produce a similar pitch angle without a great deal of tuning (as the simulations by Meru et al were not intended to reproduce the exact ALMA image, but were testing which scenarios produced morphologies that were consistent with the observations)

If Elias 2-27 is stable against GI, but undergoing encounters with an external companion, our simulations show it should produce asymmetric arms which have pitch angles that vary with radius, and a larger mean pitch overall. Due to the relatively large disc mass in both simulations (and the relatively massive companion in the Companion simulation), we find that the typical expressions for companion-driven arms (e.g. Zhu et al. 2015) are a poor fit for the spirals. We also find that other spiral functions (such as the power spiral) are generally a worse fit to the data.

The current ALMA observations indicate that Elias 2-27 does possess tightly wound, symmetric logarithmic spiral structure, suggesting that it is in fact a GI disc. However, we note that Pérez et al. (2016)’s fits to the arms assume symmetry, precluding the study of an important observable for determining the origin of spiral structure. We recommend that future observations of Elias 2-27 (and other discs with midplane-driven spiral

arms) conduct fitting of individual arms rather than a single, simultaneous fit to all arms.

4. CONCLUSIONS

In this paper, we have conducted a comparative spiral morphology study on massive protostellar disc simulations tuned to reproduce the observed spiral structure in the protostellar disc Elias 2-27 (Pérez et al. 2016). One simulation presented an isolated disc with $Q \lesssim 1$ resulting in gravitational instability (GI), the other presented a disc with $Q > 2$, undergoing perturbations during an encounter with a 10 Jupiter mass companion (Companion).

Using the TACHE algorithm (Forgan et al. 2018b) on the simulation data, we identify the spine points associated with individual spiral arms for both simulations, and fit these spine points to a variety of spiral functions. Our results show that the two simulations have markedly different “ground truth” spiral morphology, and yet produce similar “observed” morphologies under unsharp masking (Meru et al. 2017).

From this study, and from previous work, we identify several key discriminators between GI and companion-driven spiral structure for the case of Elias 2-27. The GI disc exhibits pure, symmetric logarithmic spirals of constant pitch angle, whereas the Companion disc shows asymmetric logarithmic spirals, with a pitch angle that varies with radius.

In particular, we show that asymmetry between spiral arms is a key observable, and as such observers *must* attempt to fit spirals individually, rather than assuming a single set of fit parameters for all spirals. It is worth noting that arm asymmetry is sensitive to the orbital phase of the companion, and this should be factored into any predictions based purely on asymmetry.

In conjunction with Meru et al. (2017), this study demonstrates that the current ALMA observations do not yet differentiate between the spiral structures observed in the GI and Companion simulations for Elias 2-27. This highlights the need for synthetic observations as a tool to both evaluate numerical simulations and to interpret real observations (cf Haworth et al. 2017).

In summary: we recommend further high resolution observations of Elias 2-27 to determine the source of its spiral structure. Between the two scenarios tested here, our simulations predict morphological differences on scales of 10-20au in the outer regions of the disc. Therefore, future observations of Elias 2-27 with angular resolutions of tens of milli-arcseconds have the potential to determine the source of its spiral structure. We predict that if these observations continue to show tightly

wound, symmetric spiral structure, then Elias 2-27 is indeed the first observed self-gravitating protostellar disc.

DHF gratefully acknowledges support from the ECOGAL project, grant agreement 291227, funded by the European Research Council under ERC-2011-ADG. JDI and FM acknowledge support from the DISCSIM project, grant agreement 341137 under ERC-2013-ADG. FM also acknowledges support from The Leverhulme Trust, the Isaac Newton Trust and the Royal Society Dorothy Hodgkin Fellowship. The authors warmly thank the anonymous referee for their com-

ments which helped to clarify the manuscript. This work used the Darwin DiRAC HPC cluster at the University of Cambridge, and the Cambridge COSMOS SMP system funded by ST/J005673/1, ST/H008586/1 and ST/K00333X/1 grants. This research made use of NASA's Astrophysics Data System Bibliographic Services.

Facilities: Darwin DiRAC HPC cluster, University of Cambridge

Software:

TACHE: <https://github.com/dh4gan/tache>

SCIPY: <http://www.scipy.org>

REFERENCES

- Andrews, S. M., Wilner, D. J., Hughes, A. M., Qi, C., & Dullemond, C. P. 2009, *ApJ*, 700, 1502
- Benisty, M., Juhasz, A., Boccaletti, A., et al. 2015, *A&A*, 578, L6
- Bertin, G., & Lin, C. C. 1996, *Spiral structure in galaxies a density wave theory*, ISBN0262023962
- Cossins, P., Lodato, G., & Clarke, C. J. 2009, *MNRAS*, 393, 1157
- Crida, A., Morbidelli, A., & Masset, F. 2006, *Icarus*, 181, 587
- Dong, R., Zhu, Z., Fung, J., et al. 2016, *The Astrophysical Journal Letters*, 816, L12
- Evans, M. G., Ilee, J. D., Hartquist, T. W., et al. 2017, *MNRAS*, 470, 1828
- Forgan, D., Bonnell, I., Lucas, W., & Rice, K. 2016a, *MNRAS*, 457, 2501
- Forgan, D., & Rice, K. 2013a, *MNRAS*, 432, 3168
- . 2013b, *MNRAS*, 433, 1796
- Forgan, D., Rice, K., Cossins, P., & Lodato, G. 2011, *MNRAS*, 410, 994
- Forgan, D. H., Hall, C., Meru, F., & Rice, W. K. M. 2018a, *MNRAS*, 474, 5036
- Forgan, D. H., Ilee, J. D., Cyganowski, C. J., Brogan, C. L., & Hunter, T. R. 2016b, *MNRAS*, 463, 957
- Forgan, D. H., Ramón-Fox, F. G., & Bonnell, I. A. 2018b, *ArXiv e-prints*, arXiv:1802.01364
- Fukagawa, M., Tamura, M., Itoh, Y., et al. 2006, *ApJL*, 636, L153
- Fung, J., & Dong, R. 2015, *ApJL*, 815, L21
- Galvagni, M., & Mayer, L. 2014, *MNRAS*, 437, 2909
- Gammie, C. F. 2001, *ApJ*, 553, 174
- Goodman, J., & Rafikov, R. R. 2001, *ApJ*, 552, 793
- Greaves, J. S., & Rice, W. K. M. 2010, *MNRAS*, 407, 1981
- Hall, C., Forgan, D., Rice, K., et al. 2016, *MNRAS*, 458, 306
- Hall, C., Rice, K., Dipierro, G., et al. 2018, *MNRAS*, 477, 1004
- Haworth, T. J., Glover, S. C. O., Koepferl, C. M., Bisbas, T. G., & Dale, J. E. 2017, *ArXiv e-prints*, arXiv:1711.05275
- Isella, A., Carpenter, J. M., & Sargent, A. I. 2009, *ApJ*, 701, 260
- Laughlin, G., & Bodenheimer, P. 1994, *ApJ*, 436, 335
- Lin, C. C., & Shu, F. H. 1964, *ApJ*, 140, 646
- Lin, D. N. C., & Papaloizou, J. 1986, *ApJ*, 309, 846
- Lodato, G., & Rice, W. K. M. 2005, *MNRAS*, 358, 1489
- Luhman, K. L., & Rieke, G. H. 1999, *ApJ*, 525, 440
- Meru, F. 2015, *MNRAS*, 454, 2529
- Meru, F., Juhász, A., Ilee, J. D., et al. 2017, *ApJL*, 839, L24
- Muto, T., Grady, C. A., Hashimoto, J., et al. 2012a, *ApJL*, 748, L22
- . 2012b, *ApJL*, 748, L22
- Natta, A., Testi, L., & Randich, S. 2006, *A&A*, 452, 245
- Pérez, L. M., Carpenter, J. M., Andrews, S. M., et al. 2016, *Science*, 353, 1519
- Rafikov, R. R. 2002, *ApJ*, 569, 997
- Ricci, L., Testi, L., Natta, A., & Brooks, K. J. 2010, *A&A*, 521, A66
- Rice, W. K. M., Lodato, G., & Armitage, P. J. 2005, *MNRAS*, 364, L56
- Stamatellos, D., Hubber, D. A., & Whitworth, A. P. 2007, *MNRAS*, 382, L30
- Tomida, K., Machida, M. N., Hosokawa, T., Sakurai, Y., & Lin, C. H. 2017, *ApJL*, 835, L11
- Toomre, A. 1964, *ApJ*, 139, 1217
- Whitehouse, S. C., & Bate, M. R. 2006, *MNRAS*, 367, 32
- Whitehouse, S. C., Bate, M. R., & Monaghan, J. J. 2005, *MNRAS*, 364, 1367

Zhu, Z., Dong, R., Stone, J. M., & Rafikov, R. R. 2015,
ApJ, 813, 88